Table 1 MDO Census

Galaxy	Type	D (Mpc)	$M_{B,\mathrm{bulge}}$	F_{radio} (mJy)	$\frac{L_{\mathrm{radio}}}{L_{\mathrm{radio}}(\mathrm{SgrA^*})}$	$M_\bullet$ ($\mathrm{M}_\odot$)	$\log\frac{M_\bullet}{M_{\mathrm{bulge}}}$	References
M 31	Sb	0.7	-18.82	0.03	0.3	3×10^7	-3.31	Dressler & Richstone 88; Kormendy 88c; Richstone+90 (also M 32); Bacon+94
NGC 3115	S0/	8.4	-19.90	<0.33	$<460.$	1×10^9	-1.92	Kormendy & Richstone 92
M 32	E	0.7	-15.51	$<4.$	$<54.$	2×10^6	-2.60	Tonry 84, 87; Dressler & Richstone 88; van der Marel+94b; Qian+95; Dehnen 95
NGC 4594	Sa/	9.2	-21.21	$\sim10^2$	$\sim2\times10^5$	5×10^8	-2.99	Kormendy 88d; Emsellem+94a
Galaxy	Sbc	0.0085	-17.65	7×10^2	1.	2×10^6	-3.78	McGinn+89; Sellgren+90; Kent 92; Evans & de Zeeuw 94; Haller+95
NGC 3377	E	9.9	-18.80	<1.2	$<2\times10^3$	8×10^7	-2.24	Kormendy+95b
NGC 4258	Sbc	7.5	-17.3	1.8	2×10^3	4×10^7	-2.05	Miyoshi+95
M 87	E	15.3	-21.42	$\sim10^3$	$\sim7\times10^6$	3×10^9	-2.32	Sargent+78; Harms+94; van der Marel 94a

Notes – Col. 2, morphological type (de Vaucouleurs & Pence 1978; de Vaucouleurs *et al.* 1991). Col. 3: Distance (mostly Faber *et al.* 1995; for NGC 4258: Faber 1995, personal communication). Values are only slightly modified from Lynden-Bell *et al.* (1988) and Faber *et al.* (1989). They are based on a large-scale flow field solution, an assumed Virgo Cluster distance of 15.3 Mpc, and a Hubble constant of $H_0 = 80$ km s^{-1} Mpc^{-1}. They agree well with distances in Tully *et al.* (1992). BH mass scales as D; $\log M_\bullet/M_{\rm bulge}$ is independent of D. Demographic conclusions are not sensitive to the distance scale: see versions of Figures 1, 2, 3, and 14 for $H_0 = 50$ km s^{-1} Mpc^{-1} in Kormendy (1993, 1994). Col. 4: B-band bulge absolute magnitude, based mostly on B_T in de Vaucouleurs *et al.* (1991). Exceptions are M 31 ($V_T = 3.28$ from Kent 1987b), NGC 3115 ($B_T = 9.75$ from Capaccioli *et al.* 1987), NGC 4594 ($B_T = 8.65$ from Burkhead 1986), and the Galaxy (M_B adapted from Kent *et al.* 1991). Galactic absorptions are taken from Burstein & Heiles (1984). Assumed bulge-to-total luminosity ratios are: M 31: $B/T = 0.24$ (de Vaucouleurs 1958), NGC 3115: 0.94 (Capaccioli *et al.* 1987), NGC 4594: 0.93 (Burkhead 1986), NGC 4258: 0.063 (an approximate value near the bottom of the distribution of B/T values for Sbc galaxies, see Simien & de Vaucouleurs 1986). Col. 5: Nuclear radio continuum flux. Wavelengths and sources: M 31: 3.6 cm, Crane *et al.* (1992, 1993a); NGC 3115: 6 cm, Fabbiano *et al.* (1989); M 32: 21 cm, Hummel (1980); NGC 4594: 3.6 – 21 cm, van der Kruit (1973), de Bruyn *et al.* (1976), Hummel (1980); the Galaxy: 6 cm, Zhao *et al.* (1992); NGC 3377: 6 cm, Fabbiano *et al.* (1987); NGC 4258: 6 cm, Turner & Ho (1994); M 87: 2 cm, 21 cm, Hummel (1980), Biretta *et al.* (1983). The value for M 87 depends on how much of the jet is included at the spatial resolution of the observations. Col. 6: Nuclear radio luminosity in units of the luminosity of Sgr A* at the same wavelength. Sgr A* fluxes are 0.5, 0.7, 0.8, and 0.9 Jy at 21 cm, 6 cm, 3.6 cm, and 2 cm, respectively (Zhao *et al.* 1992). Col. 7: MDO mass for the best-fitting isotropic dynamical model. Col. 8: logarithm of MDO mass fraction; $M_{\rm bulge}$ is the bulge mass determined from $M_{B,{\rm bulge}}$ and from M/L ratios measured in the BH papers (Col. 9). For NGC 4258, $M/L_r = 4.08$ is from Kent (1987a); this disk value is probably a good approximation for the bulge-like component as well. Col. 9: MDO references; *et al.* is abbreviated as + and (e. g.) 1994 as 94. Only papers that measured $M_\bullet$ are included; additional papers that confirm the kinematics are discussed in Sections 4 and 5.