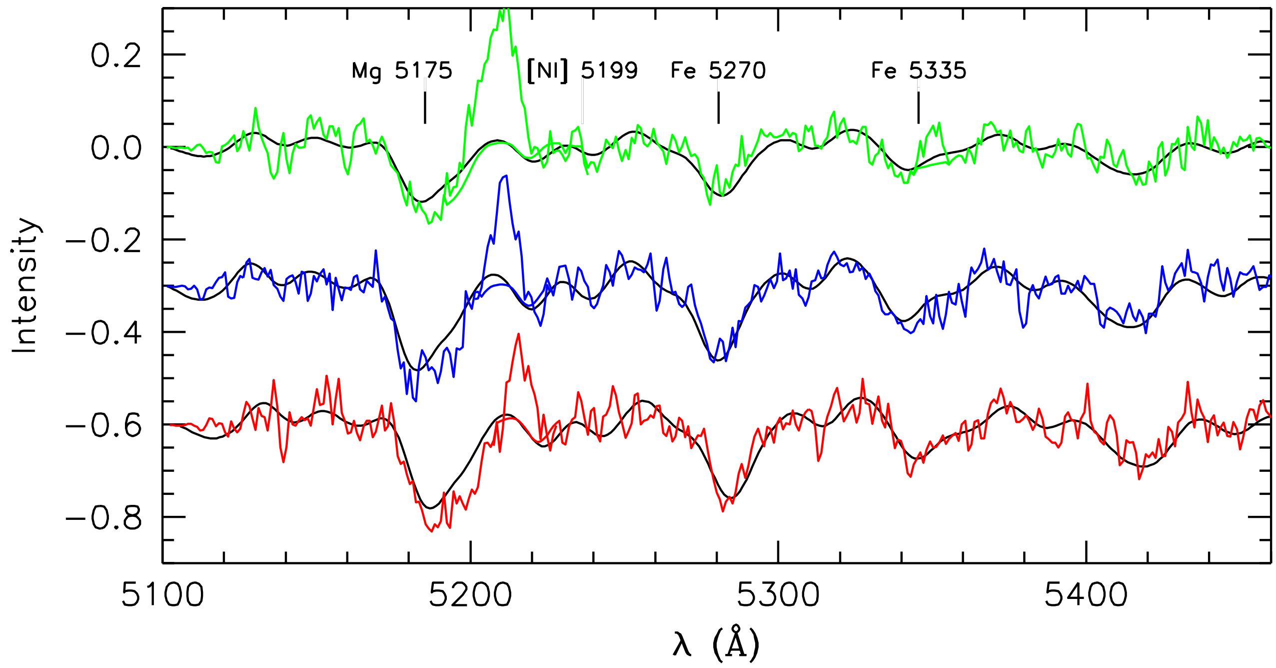

FIG. 1.—(*top*) *HST* WFPC2 images of NGC 4594: (*upper left*) Color image made from 1200 s V-band and 1050 s I-band images after 40 iterations of Lucy deconvolution. To emphasize the dust, the image sent to the blue gun of the filmwriter was an extrapolation from I through V: i.e., I_V^2/I_I, where I_V and I_I are the intensities in V and I, respectively. This panel is $9\farcs3 \times 6\farcs2$, and the displayed brightness is proportional to the square root of intensity. (*upper right*) Same image expanded by a factor of 2 ($4\farcs6 \times 3\farcs1$). (*middle right*) Same as upper right, with the $0\farcs21$ aperture positions superposed. Red indicates the receding side of the galaxy; green and blue indicate the approaching side with respect to the systemic velocity. (*middle left*) Color image $V-I$ coded so that $V-I = 1.6$ is black and $V-I = 2.3$ is white. (*bottom*) Spectra of NGC 4594 in the wavelength region used to determine the absorption-line velocities, color-coded as in the images. The spectra are shown before and after the removal of [NI] λ 5199. The black lines are the standard star spectrum broadened and velocity-shifted to fit the data. The star spectrum has been scaled using a mean line strength γ; note that the Fourier correlation quotient program is insensitive to the fact that different lines have different γ values.